Identifying planets orbiting stars other than our own might seem like an impossible task. Using even the most powerful telescopes, on Earth or in space, stars other than our own sun appear only as specs of light. Obtaining a visual image of planets orbiting these stars presents a nearly hopeless proposition. Yet, to date, astronomers (as well as novice individuals using programs available online) have identified thousands of other planets, outside our own solar system - exoplanets. How do they do it?

One method by which astronomers search for exoplanets is by measuring the extremely small changes in the light observed from a distant star as a planet transits (travels or moves across in front of) the star. This approach (the transit method) is modeled in the various scenarios in this activity.

With *Searching for Exoplanets*, you will learn:

- one of several methods by which astronomers discover planets orbiting other stars
- how the luminosity of a star is affected by the transition of a planet
- about factors used to categorize stars (luminosity, color, mass, radius)
- characteristics of planetary orbits (planet radius, eccentricity, rate of orbit)
- how astronomers predict whether a planet could be “habitable” by Earth-life criteria

**Materials**

You will need a laptop or desktop computer with internet access, and a scientific calculator.

**Getting Started**

*Searching for Exoplanets* will present you with five different solar system scenarios, each with a star of different size and luminosity, and an orbiting planet, also with a unique size, diameter, and orbital radius. You will be calculating other parameters relating to the sun-planet system and, ultimately, try to determine if you think the planet might be in what astronomers would refer to as the “habitable zone.”


Then, go to the following link to run the simulation  [https://bit.ly/39tkTP9](https://bit.ly/39tkTP9)

**Exploring the scenarios**

Each of the five (numbered 1-5) scenarios you will be running is designed to generate data for an exoplanet-star system. The data (the light curve graph) will look similar to the example in Figure 2. Below, is a view of the simulation window; highlights are identified in red.
When you are ready to begin your investigation,

- Enter a scenario of your choice, from 1-5 when prompted (or select Run from the toolbar).
- You can toggle the simulation with pause and run (or use the space bar). Pause to be able to select and read data values.
- You can also grab the view and zoom in and out of the run to re-scale the view.
- Run and work through the calculations for one scenario at a time to become familiar with the interpretation and analysis of the data you collect.
- Record your findings in the data table. (You might want to copy the table on a larger page.)

Understanding the data

The key piece of evidence you will need for your analysis is the light curve. Each dip in the graph corresponds to the transit of a planet across a star, resulting in an infinitesimally small change in brightness. Knowing the time between the transits (dips), you can figure out the period of the planet’s orbit. You can use your mouse to select points on the graph so you can more easily calculate the period and change in luminosity.

Work through each of the calculations to figure out the identifying characteristics for the planet in each of the five scenarios. In the end, you will be able to decide for yourself which of the planets in the scenarios could potentially support life!

Below are the steps you will need to complete the table for the planet in each scenario:
1) Use the light curve (the graph) from the simulation to determine the **period** of the planet by counting the number of Earth days between dips.

2) Convert the period from Earth days into **seconds** (there are 89,280 seconds per day).

3) Use the light curve (graph), to calculate $\Delta L$, the **change in brightness** that occurs when the planet transits across the star. $\Delta L$ can be calculated directly from the graph as the difference between the observed brightness of the star during a transit (example below).

4) Now, let’s calculate the **radius of the planet**.

   You will need to know the radius of the star (given on the Information Table), and your measured change in brightness.

   The **radius of the star** (stellar radius) $R_*$, is will be given in the Information Table for each scenario. (The $*$ symbol means that it relates to the star.) We will re-arrange the equation for you, below. Use the final equation to calculate the **radius of the planet** $r_p$, and record your answer in the table. (You do not need to report all of the digits your calculator gives you. Three digits will be good.)

   $\Delta L = \left( \frac{r_p}{R_*} \right)^2$  

   or,  

   \[
   \text{change in brightness} = \left( \frac{\text{radius of planet}}{\text{radius of star}} \right)^2
   \]

   First, take the square root of both sides 

   $\sqrt{\Delta L} = \frac{r_p}{R_*}$

   Then, multiply both sides by $R_*$ and calculate $r_p$ 

   $\Delta L \cdot R_* = r_p$

   (Check to see if your answers are reasonable. The values for the planet radii should be in the range of $10^6$ to $10^8$ meters.)

5) Next, we will calculate the **orbital radius**, the radius of the planetary orbit.

   Using the value for $M_*$ (mass of the star) provided in the Information Table, and the values for $P$, the **orbital period** (in seconds), that you determined from the light curve, you can calculate the **orbital radius** ($r$). Record your results in the table. (See the suggestion, that follows.) You will need the gravitational constant, $G = 6.67 \times 10^{-11} \frac{m^3}{kg \cdot s^2}$, given in the Information Table.
\[ r = \left( \frac{GM_P^2}{4\pi^2} \right)^{1/3} \quad \text{or, orbital radius} = \left( \frac{\text{Gravitational constant} \times \text{mass of star} \times (\text{period of orbit})^2}{4\pi^2} \right)^{1/3} \]

**Note:** Solving the above equation is not trivial! There are different ways to approach this and not all brands of calculators work the same. If you need a suggestion, try:

a. **Square** P (the period), you can use the \( x^2 \) function on your calculator, or just \( P \times P \).

b. Then **multiply** that answer times the **mass of the star**

c. **Multiply** that answer times the gravitational constant \( G, 6.67 \times 10^{-11} \frac{m^3}{kg \cdot s^2} \)

d. Then divide by the denominator (use parentheses). In other words, \( \frac{4 \times \pi^2}{} \).  

e. Then take that answer to the 1/3 power, in other words, \( (\text{your answer})^{1/3} \). You can

(Check to see if your answers are reasonable. The values for the orbital radii should be in the range of \( 10^9 \) to \( 10^{11} \) meters.)

6) Now, you can actually calculate the **mass of the exoplanet**!

A unique and mutual gravitational attraction exists between a planet, and the star it orbits. We will provide a value for that attraction as a “Force of Gravity” in the Information Table for each star/planet system. Using the orbital radius and Newton’s Law of gravitation you will be able to determine the **mass of the exoplanet** for each system. Again, you will need the gravitational constant, \( G = 6.67 \times 10^{-11} \frac{m^3}{kg \cdot s^2} \).

\[ m_p = \frac{F_{Gr}^2}{GM_*} \quad \text{or, mass of planet} = \frac{\text{Force of gravity} \times (\text{orbital radius})^2}{\text{Gravitational constant} \times \text{mass of star}} \]

**Note:** There are multiple ways to solve an equation like this, but if you have trouble, try this:

a. **Square the radius** (use the \( x^y \) function on your calculator), or just do radius*radius.

b. **Multiply** the answer by **force of gravity** (from the data table).

c. **Divide** that answer by the gravitational constant, \( 6.67 \times 10^{-11} \frac{m^3}{kg \cdot s^2} \)

d. **Divide** that answer by the **stellar mass** (mass of the star).

7) (Check to see if your answers are reasonable. The values for the masses of the exoplanets should be in the range of \( 10^{24} \) to \( 10^{27} \) kilograms.)
8) Could the planet be habitable? Can you make a prediction?

The mass of an exoplanet can give clues about whether the planet might be able to support life. A life-supporting planet must be massive enough to maintain an atmosphere. But having an atmosphere alone is not sufficient to sustain life.

The most important condition for a planet to be habitable is that the planet must contain liquid water. Detecting the presence of liquid water is exceptionally difficult but one simple test can be applied to exoplanets. A planet must be the correct temperature to possibly contain liquid water. The effective temperature of a planet is determined by the brightness of the host star and the planet’s orbital radius.

Using the given value for the luminosity, \( L \) listed in the data table, and the value of the orbital radius, \( r \), that you calculated in Step 5, you can now calculate the effective temperature, \( T \), of the planet using the equation below.

**Note:** You will use the luminosity, \( L \), from the data table, not the change in brightness, \( \Delta L \), that you calculated. The \( 1.4 \times 10^6 \) term is a proportionality constant, relative to the calculation.

\[
T = \left(1.4 \times 10^6 + \frac{L}{r^2}\right)^{1/4} \quad \text{or} \quad \text{temperature} = \left(1.4 \times 10^6 + \frac{\text{luminosity}}{(\text{orbital radius})^2}\right)^{1/4}
\]

**Note:** Again, these steps may help you if you would like a suggestion:

a. Enter the value for the luminosity of the star (from the Information Table).

b. Divide by the orbital radius, and then divide by the orbital radius again.

c. Then take that answer to the \( 1/4 \) power, in other words, \((\text{your answer})^{1/4}\). You can

Water freezes at 273 K and water boils at 373 K. To a rough estimate, an exoplanet is within the habitable zone of its solar system if the effective temperature is between 273 — 373 K.
<table>
<thead>
<tr>
<th>Scenario</th>
<th>Stellar Radius (meters)</th>
<th>Stellar Mass (kilogram)</th>
<th>Force of Gravity (Newtons)</th>
<th>Luminosity of the Star (Watts)</th>
<th>Gravitational Constant, $G \ m^3/kg \cdot s^2$</th>
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<th>Scenario</th>
<th>Orbital Period (Earth Days)</th>
<th>Orbital Period (seconds)</th>
<th>$\Delta L_r$, Change in Brightness (Watts)</th>
<th>Planet Radius (meters)</th>
<th>Orbital Radius (meters)</th>
<th>Planet Mass (kilograms)</th>
<th>Planet Temperature (Kelvin)</th>
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Searching for Exoplanets - Saturday Morning Astrophysics at Purdue; Dr. David Sederberg, dsederbe@purdue.edu
Additional Resources

1. https://exoplanets.nasa.gov/alien-worlds/historic-timeline/#first-exoplanet-found-around-a-main-sequence-star
3. https://www.planethunters.org/